

The European Large Area ISO Survey - VII. ROSAT observations of ELAIS sources

Article (Published Version)

Basilakos, S, Georgantopoulos, I, Pérez-Fournon, I, Efstathiou, A, Rowan-Robinson, M, Cabrera-Guerra, F, González-Solares, E, Alexander, D M, Serjeant, S and Oliver, S (2002) The European Large Area ISO Survey - VII. ROSAT observations of ELAIS sources. *Monthly Notices of the Royal Astronomical Society*, 331 (2). pp. 417-422. ISSN 0035-8711

This version is available from Sussex Research Online: <http://sro.sussex.ac.uk/20841/>

This document is made available in accordance with publisher policies and may differ from the published version or from the version of record. If you wish to cite this item you are advised to consult the publisher's version. Please see the URL above for details on accessing the published version.

Copyright and reuse:

Sussex Research Online is a digital repository of the research output of the University.

Copyright and all moral rights to the version of the paper presented here belong to the individual author(s) and/or other copyright owners. To the extent reasonable and practicable, the material made available in SRO has been checked for eligibility before being made available.

Copies of full text items generally can be reproduced, displayed or performed and given to third parties in any format or medium for personal research or study, educational, or not-for-profit purposes without prior permission or charge, provided that the authors, title and full bibliographic details are credited, a hyperlink and/or URL is given for the original metadata page and the content is not changed in any way.

The European Large Area *ISO* Survey – VII. *ROSAT* observations of ELAIS sources

S. Basilakos,¹ I. Georgantopoulos,^{2*} I. Pérez-Fournon,^{3,4} A. Efstathiou,¹
M. Rowan-Robinson,¹ F. Cabrera-Guerra,³ E. González-Solares,³ D. M. Alexander,⁵
S. Serjeant⁶ and S. Oliver⁷

¹*Astrophysics Group, Imperial College London, Blackett Laboratory, Prince Consort Road, London SW7 2BW*

²*Institute of Astronomy and Astrophysics, National Observatory of Athens, Palaia Penteli, 15236, Athens, Greece*

³*Instituto de Astrofísica de Canarias, Via Lactea s/n, 38200 La Laguna, Tenerife, Spain*

⁴*Instituto de Astrofísica de Canarias, Universidad La Laguna, Tenerife, Spain*

⁵*Department of Astronomy and Astrophysics, 525 Davey Laboratory, Pennsylvania State University, University Park, PA 16802, USA*

⁶*Unit for Space Sciences and Astrophysics School of Physical Sciences, University of Kent, Canterbury, Kent CT2 7NR*

⁷*Astronomy Centre, CPES, University of Sussex, Falmer, Brighton BN1 9QJ*

Accepted 2001 November 22. Received 2001 November 21; in original form 2001 September 14

ABSTRACT

We present a cross-correlation between the European Large Area *ISO* Survey (ELAIS) 15- μm *ISO* survey with the *ROSAT* (0.1–2 keV) all-sky survey and the pointed observations WGACAT source catalogue. The resulting sample contains 15 objects. Optical spectroscopic identifications exist for 13 objects: six broad-line quasi-stellar objects (QSOs), four narrow-line galaxies (NLGs) or type 2 active galactic nuclei (AGN) and three stars. We have used both the X-ray to infrared (IR) luminosity flux ratio f_x/f_{IR} and the X-ray hardness ratios diagnostics to estimate the amount of obscuration in these objects. The X-ray spectrum of the narrow-line galaxies does not present strong evidence for obscuration; however, the low f_x/f_{IR} ratio combined with the high X-ray luminosities suggest that at least one of the NLGs is associated with an obscured Seyfert nucleus. Four out of the six QSOs present high f_x/f_{IR} ratios and steep X-ray spectra with $\Gamma > 2$. One QSO (ELAISC15–J133442+375736) at a redshift of $z = 1.89$, has an abnormally low X-ray/IR flux ratio, with its infrared luminosity approaching that of an hyperluminous galaxy ($\sim 10^{12.98} h^{-2} L_{\odot}$). Finally, one radio-loud QSO is the hardest X-ray source in our sample, presenting strong evidence for a high absorbing column ($N_{\text{H}} \sim 10^{22} \text{ cm}^{-2}$).

Key words: cosmology: observations – infrared: galaxies – X-rays: galaxies.

1 INTRODUCTION

The standard Unification model (Antonucci & Miller 1985) asserts that the nuclei in both type 1 and type 2 active galactic nuclei (AGN) have basically the same structure containing a super-massive black hole, an accretion disc, a molecular torus and a broad line region. Then their classification as type 1 or type 2 AGN depends solely on the viewing angle. Specifically, if the source is observed at a sufficiently high inclination angle and thus the line of sight intersects the torus, it would be classified as a Seyfert 2, whereas for all other orientations it would be deemed to be a Seyfert 1. X-ray observations of Seyfert-2 galaxies with *Ginga*, *ASCA* and *RXTE* (Smith & Done 1996; Turner et al. 1997; Georgantopoulos & Papadakis 2001) observe large absorbing

columns supporting this scenario. However, optical surveys for AGN have failed so far to produce large numbers of obscured AGN beyond the local Universe.

In contrast, recent X-ray surveys have proved very useful in finding evidence for the presence of such obscured AGN population at moderate to high redshifts ($z \geq 0.1$). In particular, *ROSAT* surveys (0.1–2 keV) have detected a number of narrow-line galaxies (NLGs) the majority of which are associated with obscured AGN (Boyle et al. 1995; Schmidt et al. 1998; Lehmann et al. 2000). *ASCA* and *BeppoSAX* surveys have also found some examples of such an obscured AGN population (Boyle et al. 1998; Georgantopoulos et al. 1999; Fiore et al. 1999; Akiyama et al. 2000), in the hard 2–10 keV band, which is less prone to photo-electric absorption. These AGN present column densities typically higher than 10^{23} cm^{-2} while in some cases their optical spectra may present broad lines with only moderate optical reddening. The

*E-mail: ig@astro.noa.gr

above results have been corroborated by recent *Chandra* and *XMM* surveys (e.g. Mushotzky et al. 2000; Brandt et al. 2001; Hasinger et al. 2001). Moreover, *ASCA* (Nakanishi et al. 2000), *Chandra* (Norman et al. 2001) and *XMM* (Lehmann et al. 2001) have produced the first examples of the long-sought type 2 quasi-stellar object (QSO) population, i.e. with no signs of broad emission lines, similar to Seyfert 2 galaxies in the local Universe.

Obscured AGN should emit copious amounts of infrared (IR) radiation as the obscuring matter reprocesses the optical radiation. Therefore, the combination of X-ray and IR observations is a powerful tool in the detection of such obscured objects. In particular, using both the f_x/f_{IR} and the X-ray hardness ratios, we can detect the obscured AGN, due to the fact that the latter quantities are excellent indicators of high photoelectric absorption.

Green, Anderson & Ward (1992) combined far-infrared (from *IRAS*) and X-ray (from *Einstein*) data, finding a significant correlation between luminosities in the 60- μm and 0.5–4.5 keV. They also find that the f_x/f_{IR} ratios of broad-line AGN are significantly higher than those of narrow-line AGN and star-forming galaxies. In a similar way, Boller et al. (1992) cross-correlated the *ROSAT* with the *IRAS* all-sky surveys. The resulting sample consists of about 200 objects of which many are obscured Seyfert galaxies (Moran, Halpern & Helfand 1996). More recently, Gunn et al. (2001) have observed 15 sources (QSOs and narrow-line AGN), detected in deep *ROSAT* fields, with ISOPHOT onboard the *Infrared Space Observatory (ISO)*. The large number of narrow line objects detected suggests that large amounts of obscuring matter are present in these objects. Finally, Alexander et al. (2001) observed an area of 0.7 deg² from the ELAIS *ISO* survey (Oliver et al. 2000; Serjeant et al. 2000) to a flux limit of 10^{-13} erg cm⁻² s⁻¹ in the 2–10 keV band with *BeppoSAX*. They find 17 common sources in the hard 2–10 keV band and the 15- μm band. Surprisingly, no obscured AGN are among these sources.

Here we cross-correlate the ELAIS (Oliver et al. 2000; Serjeant et al. 2000) 15 μm *ISO* survey with the *ROSAT* all-sky survey and the pointed observations WGACAT source catalogues. Willott et al. (2001) present preliminary results on *Chandra* 0.3–10 keV observations of ACIS-I fields, covering 0.14 deg² and probing fluxes as faint as $\sim 10^{-16}$ erg cm⁻² s⁻¹ in the 2–10 keV band. The *ROSAT* observations presented here are complementary as they cover the *full* area of the ELAIS survey, albeit at much brighter fluxes. Our aim is to find a number of obscured AGN especially at high redshift and explore their nature. We note that although the *ROSAT* passband is soft (0.1–2 keV) and therefore more susceptible to photoelectric absorption as compared to either *BeppoSAX* or *Chandra*, at high redshifts the *K*-correction diminishes the *effective* obscuring column as it moves the photoelectric absorption cut-off towards lower energies.

The plan of this paper is the following: in Section 2, we describe the ELAIS and *ROSAT* catalogues used as well as the cross-correlation results, while in Sections 3 and 4 we present the discussion and conclusions.

2 OBSERVATIONS

2.1 The European Large Area *ISO* Survey (ELAIS)

ELAIS has surveyed ~ 12 deg² of the sky at 15 and 90 μm and subsets of this area at 6.75 and 175 μm using the ISOCAM (Cesarsky et al. 1996) onboard the *ISO*, (Kessler et al. 1996). The catalogue, a reliable subset of the preliminary analysis catalogue (Serjeant et al. 2000), contains 484 sources down to a flux limit of

~ 4 mJy in the 15- μm band. The positional error of the 15- μm catalogue is typically 3 arcsec (Serjeant et al. 2000). A large number of the 15- μm sources (with optical counterparts down to $R \sim 20.5$) have been spectroscopically identified containing a large fraction of AGN and star-forming galaxies (Cabrera-Guerra et al., in preparation; Pérez-Fournon et al., in preparation; Gruppioni et al., in preparation; González-Solares et al. 2001).

2.2 *ROSAT* data

We have used *ROSAT* data from the Position Sensitive Proportional Catalogue Counter (PSPC) operating in the 0.1–2.4 keV band. In particular, we have used the *ROSAT* All-Sky Survey Bright Source Catalogue (RASSBSC), *ROSAT* All-Sky Survey Faint Source Catalogue (RASSFSC) and WGACAT point source catalogues. RASSBSC is derived from the all-sky *ROSAT* survey and contains 18 811 sources (Voges et al. 1999) in the energy band 0.1–2.4 keV. RASSFSC is an extension of the RASSBSC (Voges et al. 2000) and contains approximately 106 000 sources. WGACAT is a point source catalogue generated from all *ROSAT* PSPC pointed observations (see White, Giommi & Angelini 1994). The last version of this catalogue contains about 88 000 detections, with more than 84 000 individual sources, obtained from 4160 sequences.

The PSPC has an energy resolution of 0.5 keV at 1 keV, but limited spatial resolution [full width at half maximum (FWHM) ~ 30 arcsec]. The positional rms error of the PSPC detector is typically 15 arcsec, although the exact error depends on the brightness of the source. Throughout this paper, X-ray luminosities are calculated from X-ray fluxes for a Hubble constant of $H_0 = 65$ km s⁻¹ Mpc⁻¹ and $\Omega_0 = 1$, assuming an average spectrum of $\Gamma = 2$.

2.3 The ELAIS/*ROSAT* compilation

We have performed a positional cross-correlation of the ELAIS 15- μm catalogue with the above four *ROSAT* catalogues, within $\delta\theta \leq 0.5$ arcmin. The resulting sample contains 15 objects. Note that ELAISC15-J050226-304113 and ELAISC15-J05228-304140 are associated with the same X-ray source; both ELAIS sources are located at the same redshift ($z = 0.191$). We have performed 20 simulations in order to assess the probability of chance coincidences. We offset the actual ELAIS source coordinates by a few arcmin and then we repeat the cross-correlation. We find that the chance coincidence probability is low. In particular we find that for the RASSBSC we expect < 0.05 false coincidences. For the RASSFSC and WGACAT the numbers are 0.3 and 0.7, respectively. In Fig. 1 we plot the IR flux distribution of our

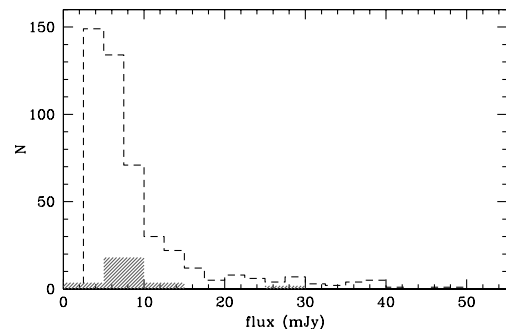


Figure 1. The flux distribution of our 15 sources relative to the ELAIS 15- μm flux distribution.

Table 1. The ELAIS/ROSAT cross-correlation. Column 8 lists the catalogues in which the sources were detected – the numbers correspond to: 1 – RASSBSC; 2 – RASSFSC; 3 – WGACAT; and 4 – ROSHRI.

Name	X-ray position	$\delta\theta$ (arcmin)	f_{IR}	f_{rad}	z	Classification	ROSAT catalogue
ELAISC15-J143143+330131	14 31 43.4+33 01 31	0.083	4.95	–	–		1,3
ELAISC15-J133451+374616	13 34 51.3+37 46 19	0.078	12.13	–	–	Star ^d	1,3,4
ELAISC15-J160623+540555	16 06 24.1+54 05 59	0.174	8.00	168	0.875	QSO ^b	2
ELAISC15-J160706+550335	16 07 04.9+55 03 58	0.414	6.03	–	–	Star ^b	2
ELAISC15-J161521+543147	16 15 20.5+54 32 12	0.423	5.74	–	0.474	QSO ^b	2,3
ELAISC15-J163310+405644	16 33 08.9+40 56 32	0.289	6.26	–	0.136	NLG ^b	2
ELAISC15-J163709+414030	16 37 08.9+41 40 54	0.401	12.51	8.37	0.765	QSO ^{b,d}	2,3
ELAISC15-J003015–430333	00 30 14.5–43 03 23	0.190	6.15	–	1.564	QSO ^a	2
ELAISC15-J003515–433355	00 35 15.9–43 33 50	0.183	29.40	–	–		2
ELAISC15-J050212–302828	05 02 12.0–30 28 23	0.083	9.38	–	0.86	QSO ^a	2,3
ELAISC15-J050226–304113	05 02 26.4–30 41 28	0.264	9.64	–	0.191	NLG ^a	3
ELAISC15-J050228–304140	05 02 26.4–30 41 28	0.398	5.02	–	0.191	NLG	3
ELAISC15-J133401+374912	13 33 59.6+37 49 12	0.28	7.77	–	0.062	NLG ^d	3,4
ELAISC15-J133414+375133	13 34 14.4+37 51 34	0.08	9.58	–	–	Star ^c	3,4
ELAISC15-J133442+375736	13 34 44.9+37 57 17	0.500	4.13	–	1.89	QSO ^c	3

Classifications are obtained from: ^a Grupponi et al. (in preparation); ^b Cabrera-Guerra et al. (in preparation) or Pérez-Fournon et al. (in preparation); ^c McHardy et al. (1998); ^d NASA Extragalactic Database (NED).

Table 2. The X-ray properties.

Name	$\log L_x$ erg s ⁻¹	f_x $\times 10^{-13}$ erg cm ⁻² s ⁻¹	Γ
ELAISC15-J143143+330131		9.80	3.17 ± 0.15
ELAISC15-J133451+374616		9.95	2.22 ± 0.07
ELAISC15-J160623+540555	44.91	3.15	0.43 ± 0.78
ELAISC15-J160706+550335		1.35	2.84 ± 0.81
ELAISC15-J161521+543147	44.44	4.07	2.62 ± 0.41
ELAISC15-J163310+405644	42.65	0.90	0.67 ± 1.52
ELAISC15-J163709+414030	45.11	6.77	2.84 ± 0.22
ELAISC15-J003015–430333	45.44	3.00	–
ELAISC15-J003515–433355		5.28	–
ELAISC15-J050212–302828	44.98	3.90	2.51 ± 0.33
ELAISC15-J050226–304113	43.13	1.37	1.85 ± 0.48
ELAISC15-J050228–304140	43.13	1.37	1.85 ± 0.48
ELAISC15-J133401+374912	41.28	0.19	2.11 ± 0.48
ELAISC15-J133414+375133		0.16	3.95 ± 0.48
ELAISC15-J133442+375736	44.32	0.15	3.28 ± 0.48

sources relative to the full ELAIS sample. For 13 of our sources, spectroscopic identifications exist either through the ELAIS spectroscopic follow-up program or through the literature. Details of the 15 sources are presented in Table 1 as follows: column (1) ELAIS name; column (2) X-ray position; column (3) angular separation between the IR and X-ray position; column (4) the 15- μm flux (mJy) as listed in Serjeant et al. (2000); column (5) radio flux (mJy) as listed in Ciliegi et al. (1999); column (6) redshift; column (7) object classification (the sources of the classifications are given in the table footnotes); column (8) the ROSAT catalogues in which the source has been detected. Note that three of our sources have been also detected in the ROSAT High-Resolution Imager (HRI) source catalogue, ROSHRI, (<http://ledas-www.star.le.ac.uk/rosat/rra/roshri>) and therefore have more accurate X-ray positions. Hence, we list here the HRI X-ray positions for these sources. In Table 2 we present the X-ray properties of our sources: column (2) gives the logarithmic X-ray luminosity (erg s⁻¹) in the 0.1–2 keV band; column (3) the X-ray flux (0.1–2 keV) in units of 10⁻¹³ erg s⁻¹ cm⁻² (note that a few of our objects have been detected in more than one ROSAT catalogue; for these we present here the X-ray flux from the observation with the highest photon statistics); column (4) X-ray photon index together with the 1 σ

error. The photon indices were estimated from the hardness ratio defined as $(h - m)/(h + m)$ where h and m correspond to the 0.9–2 and 0.5–0.9 keV (RASSFSC and RASSBSC) and the 0.9–2 keV and 0.4–0.9 bands (WGACAT). Unfortunately, two objects have very poor photon statistics and therefore their spectral index could not be calculated (ELAISC15-J003015-430333 and ELAISC15-J003515-433355).

2.4 Notes on individual objects

(i) ELAISC15–J160623+540555. This is the hardest object in X-rays, having $\Gamma = 0.43 \pm 0.78$. Recently, Pérez-Fournon et al. (in preparation) discovered that this object presents broad lines. The redshift as measured from the Mg II line is 0.875. If we assume that $\Gamma = 1.9$, then the column density is $N_{\text{H}} \sim 10^{22}$ cm⁻² at the rest-frame of the QSO. It is also detected in the radio follow-up of the ELAIS fields with flux of 167 mJy (Ciliegi et al. 1999). Furthermore, its f_x/f_{IR} ratio $\log(f_x/f_{\text{IR}}) = -5.17$ is comparable to those of QSOs.

(ii) ELAISC15–J143143+330131. This unidentified source presents a steep X-ray spectrum $\Gamma = 3.17 \pm 0.15$ together with a

high f_x/f_{IR} ratio, $\log(f_x/f_{\text{IR}}) = -4.47$. The above suggest that most probably this source is a QSO.

(iii) ELAISC15–J160706+550335. Again, for this source we have no optical identification. The $\log(f_x/f_{\text{IR}}) = -5.41$ ratio for this object places it marginally in the narrow-line galaxy regime. Still, it has a steep spectrum with $\Gamma = 2.84 \pm 0.81$. If this object is associated with an obscured AGN, the steep X-ray emission may be due to a scattered component.

(iv) ELAISC15–J133442+375736. This QSO ($z = 1.89$) presents a very steep spectrum in X-rays ($\Gamma = 3.28 \pm 0.48$) with X-ray luminosity of $L_x = 2 \times 10^{44} \text{ erg s}^{-1}$. This object could be associated with an intrinsically weak X-ray source. Interestingly, the infrared luminosity is very high, $L_{\text{IR}} \sim 10^{12.98} h^{-2} L_{\odot}$, approaching the luminosity of a hyperluminous galaxy.

(v) ELAISC15–J163310+405644. This is classified as a star-forming galaxy according to Pérez-Fournon et al. (in preparation). Its luminosity is $L_x = 4 \times 10^{42} \text{ erg s}^{-1}$, higher than those of known star-forming galaxies. The high X-ray luminosity is in stark contrast with the optical spectrum. This object bears close resemblance to the ‘composite’ Seyfert/star-forming galaxies of Moran et al. (1996). Unfortunately, the X-ray spectrum has poor photon statistics and thus it cannot provide additional constraints on the nature of this object. The $\log(f_x/f_{\text{IR}}) = -5.6$ is low, typical of obscured Seyfert galaxies in the local Universe. Finally, considering that this source is an AGN, the estimated infrared luminosity, $L_{\text{IR}} \approx 2 \times 10^{11} h^{-2} L_{\odot}$, respectively.

(vi) ELAISC15–160706+550335. This is classified as a star, according to Pérez-Fournon et al. (in preparation). The optical magnitude is very bright ($R \sim 12$).

(vii) ELAISC15–J133451+374616. This is classified as a star according to the NASA Extragalactic Database (NED). Still, its f_x/f_{IR} ratio is high for a normal star (see Stocke et al. 1991), suggesting that this is probably an X-ray binary.

3 DISCUSSION

In Fig. 2 we plot the estimated $\log(f_x/f_{\text{IR}})$ as a function of redshift for our 15 sources. The lines model the expected evolution of the f_x/f_{IR} ratio for various classes of extragalactic object, namely QSOs, Seyfert-2 and narrow-line galaxies. The solid line (QSOs) is based on the average QSO spectral energy distribution (SED) produced by Elvis et al. (1994). The dashed line (star-forming

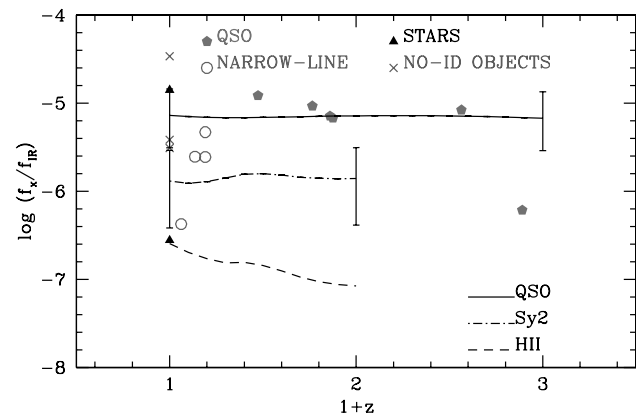


Figure 2. The X-ray/IR flux ratios for various types of galaxies as a function of $1+z$. The lines denote the predicted ratios (see Elvis et al. 1994; Alexander et al. 2001), while the errors correspond to the statistical spread in X-ray/IR colours for source type.

galaxies) and the dot-dashed line (Seyfert-2) are based on infrared SEDs generated using the Xu et al. (1998) model (see also Alexander et al. 2001). The errors on the above lines correspond to the statistical spread in X-ray/IR colours. Note that the QSO X-ray/IR ratios in our sample are, in general, in good agreement with the corresponding f_x/f_{IR} model. The only exception is ELAISC15–J133442+375736, the highest-redshift QSO ($z = 1.89$) in our sample. The narrow line galaxies are again in good agreement with the predicted f_x/f_{IR} ratio with the possible exception of ELAISC15–J050228–304140 which has a somewhat higher f_x/f_{IR} approaching the QSO regime. Note, however, that this is the ‘double’ ISO source at $z = 0.191$ which is associated with a single X-ray source. Therefore we are unable to assess the exact amount of the X-ray emission which originates from each galaxy and to draw conclusions on the nature of this object based on the f_x/f_{IR} ratio alone.

The photon indices provide additional evidence on the nature of these objects. In Fig. 3 we present the derived X-ray spectrum, Γ , as a function of redshift, while in Fig. 4 we show the $\log(f_x/f_{\text{IR}})$ ratio as a function of Γ . From Fig. 4 it becomes obvious that most QSOs have soft X-ray spectra ($\Gamma > 2$). These values are quite typical of QSO spectra in the soft *ROSAT* band (Schartel et al. 1996; Fiore, Matt & Nicastro 1997; Blair et al. 2000). Interestingly, the high-redshift QSO ELAISC15–J133442+375736, which presents the lowest f_x/f_{IR} ratio compared to the other QSOs, also shows a steep spectrum. The steep spectral index observed ($\Gamma \approx 3.3 \pm 0.5$) clearly argues against an absorbed AGN scenario. Therefore, the low f_x/f_{IR} ratio could be attributed to either an

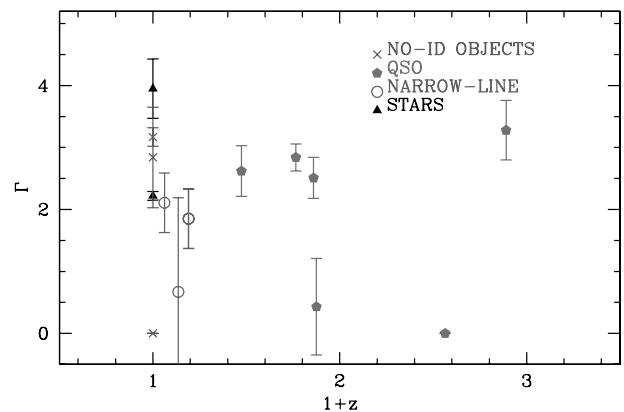


Figure 3. The X-ray spectrum as a function of $1+z$.

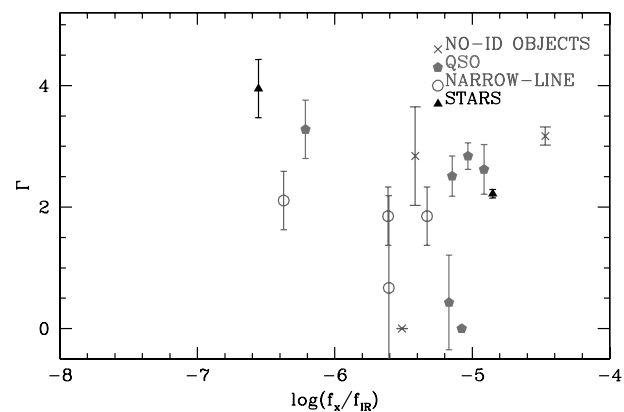


Figure 4. The X-ray spectrum versus $\log(f_x/f_{\text{IR}})$.

intrinsically low X-ray emission (similar to PG1011-040; Gallagher et al. 2001) or alternatively to high IR luminosity. Indeed, this source has a luminosity of $\sim 10^{12.98} h^{-2} L_{\odot}$. For comparison we note that our source has an f_x/f_{IR} ratio a few times lower than that of the hyperluminous source ($\sim 1.2 \times 10^{13} h^{-2} L_{\odot}$) ELAISP90–J164010+410502 (Morel et al. 2001), which has been detected in the 90- μm ELAIS sample. The hardest X-ray source in our sample is ELAISC15–J160623+540555. This is a radio-loud QSO at a redshift of $z = 0.875$. The hardness ratio corresponds to a photon index of 0.43 or alternatively to a column density of $N_{\text{H}} \sim 10^{22} \text{ cm}^{-2}$ (assuming $\Gamma = 1.9$) at the QSO's rest-frame. This is not unphysical for radio-loud QSOs. Indeed, Reeves & Turner (2000) using *ASCA* observations of a sample of 35 radio-loud QSOs, find that a large fraction of these present absorbing columns of the order of 10^{22} cm^{-2} or higher. The same conclusions have been reached by Fiore et al. (1998) using *ROSAT*.

The narrow-line galaxies have soft X-ray spectra all consistent with $\Gamma \sim 2$. Only ELAISC15–J163310+405644 presents a harder X-ray spectrum ($\Gamma = 0.67 \pm 1.52$) but within the large uncertainty this is still consistent with the spectra of the other narrow-line objects. Such soft X-ray spectra are compatible with X-ray observations of star-forming galaxies in the *ROSAT* band (see Read, Ponman & Strickland 1997; Dahlem, Weaver & Heckman 1998). The X-ray luminosities of these objects yield additional clues on the nature of these objects. It is becoming clear that star-forming galaxies cannot exceed luminosities of $10^{42} \text{ erg s}^{-1}$ in the X-ray band (e.g. Moran et al. 1996; Zezas, Georgantopoulos & Ward 1998). Then the high luminosity of ELAISC15–J163310+405644 implies the presence of an active nucleus in this object. Still, the optical spectrum presents no sign of nuclear activity. This object has very similar properties to the ‘composite’ galaxies presented in detail in Veron, Goncalves & Veron-Cetty (1997) and Moran et al. (1996). The ‘composites’ have optical spectra which would classify these as star-forming galaxies on the basis of the diagnostic line ratio diagram of Veilleux & Osterbrock (1987), but still their X-ray luminosities are typical of AGN. The ‘double’ *ISO* source again presents a high X-ray luminosity ($\sim 10^{43} \text{ erg s}^{-1}$). Then both or at least one of the sources are associated with an AGN. Alternatively, the X-ray source could be associated with a cluster of galaxies (indeed both ELAIS sources have the same redshift) while the IR emission could originate from star-forming processes in the two galaxies. Finally, ELAISC15–J133401+374912 presents a low X-ray luminosity typical of star-forming galaxies suggesting that this may be the only ‘bona-fide’ normal galaxy in our sample. Of course the possibility that this hosts a low luminosity AGN cannot be ruled out.

4 SUMMARY AND CONCLUSIONS

We have presented a cross-correlation between the ELAIS (Oliver et al. 2000; Serjeant et al. 2000) 15- μm *ISO* survey with the *ROSAT* (0.1–2 keV) all-sky survey and pointed observations source catalogues. We found 15 common objects, 13 of which have optical identifications: six broad-line QSOs ($z = 0.47–1.89$), four NLGs or type 2 AGN ($z = 0.06–0.19$) and three stars. Utilizing the measured f_x/f_{IR} and the X-ray hardness ratios we can investigate whether our sources present large amounts of obscuration. We find that 5/6 QSOs present steep spectra, arguing against the presence of large (more than a few times 10^{21} cm^{-2}) obscuring columns. The hardest source in our sample is a QSO which has a column density of $N_{\text{H}} \sim 10^{22}$. This is a radio-loud QSO and therefore the presence of a large obscuring column is not surprising. The

highest-redshift QSO has a low f_x/f_{IR} ratio, which suggests that its X-ray emission is weak relative to its powerful IR luminosity which is borderline to that of hyperluminous galaxies.

On the other hand, the narrow-line galaxies in general have soft X-ray spectra consistent with $\Gamma \sim 2$ and also therefore with both the star-forming galaxy and AGN spectra. ELAISC15–J163310+405644 presents a harder X-ray spectrum ($\Gamma = 0.67 \pm 1.52$), but within the large uncertainty this is still consistent with the spectra of the other narrow-line objects. On the basis of the high X-ray luminosity ($> 10^{42} \text{ erg s}^{-1}$), it becomes clear that at least one of the four narrow-line objects (ELAISC15–J163310+405644) harbours an AGN. Then the low X-ray to IR flux ratio implies that we are only viewing the scattered or star-forming component in the *ROSAT* passband. This object presents additional interest because the optical spectrum classifies it as a star-forming galaxy and demonstrates that X-ray observations play a critical role in identifying the true nature of IR sources.

In conclusion, the majority of extragalactic objects in our sample falls roughly into two categories: high-redshift non-obscured radio-quiet QSOs and low-redshift absorbed AGN or star-forming galaxies. We have not found any evidence for the presence of absorbed radio-quiet QSOs at high redshift similar to those found in *BeppoSAX* or *ASCA* surveys. Deeper X-ray observations of a few of the ELAIS fields with *Chandra* and *XMM* are expected to shed more light on the number density and the nature of these enigmatic objects.

ACKNOWLEDGMENTS

We are grateful to F. La Franca and C. Gruppioni for communicating to us optical identifications for some of the sources prior to publication. This work was supported by EC Network programme ‘POE’ (grant number HPRN-CT-2000-00138). IPF, FCG and EGS were supported by project PB98-0409-CO2-01 of the Spanish Ministerio de Ciencia y Tecnología.

REFERENCES

Akiyama M. et al., 2000, *ApJ*, 532, 700
 Alexander M. D. et al., 2001, *ApJ*, 554, 18
 Antonucci R., Miller J. S., 1985, *ApJ*, 297, 621
 Blair A. J., Stewart G. C., Georgantopoulos I., Boyle B. J., Griffiths R. E., Shanks T., Almaini O., 2000, *MNRAS*, 314, 138
 Boller T., Meurs E. J. A., Brinkmann W., Fink H., Zimmermann U., Adorf H. M., 1992, *A&A*, 261, 57
 Boyle B. J., McMahon R. G., Wilkes B. J., Elvis M., 1995, *MNRAS*, 272, 462
 Boyle B. J., Almaini O., Georgantopoulos I., Blair A. J., Stewart G. C., Griffiths R. E., Shanks T., Gunn K. F., 1998, *MNRAS*, 297, L53
 Brandt W. N., Hornschemeier A. E., Schneider D. P., Alexander D. M., Bauer F. E., Garmire G. P., Vignali C., 2001, *ApJ*, 558, L5
 Cesarsky C. J. et al., 1996, *A&A*, 315, L32
 Ciliegi P. et al., 1999, *MNRAS*, 302, 222
 Dahlem M., Weaver K. A., Heckman T. M., 1998, *ApJS*, 118, 401
 Elvis M. et al., 1994, *ApJS*, 95, 1
 Fiore F., Matt G., Nicastro F., 1997, *MNRAS*, 284, 731
 Fiore F., Elvis M., Giommi P., Padovani P., 1998, *ApJ*, 492, 79
 Fiore F., La Franca F., Giommi P., Elvis M., Matt G., Comastri A., Molendi S., Gioia I., 1999, *MNRAS*, 306, L55
 Gallagher S. C., Brandt W. N., Sambruna R. M., Mathur S., Yamasaki N., 2001, *ApJ*, 546, 795
 Georgantopoulos I., Almaini O., Shanks T., Stewart G. C., Griffiths R. E., Boyle B. J., Gunn K. F., 1999, *MNRAS*, 305, 125

- Georgantopoulos I., Papadakis I. E., 2001, MNRAS, 305, 125
 González-Solares E. A. et al., 2001, in Pilbratt G. L. et al., eds, Proc. ESA Conf. ESA-SP460, The Herschel Observatory. ESA Publications Division, Noordwijk, p. 421
 Green P. J., Anderson S. F., Ward M. J., 1992, MNRAS, 254, 30
 Gunn K. F. et al., 2001, MNRAS, submitted
 Hasinger G. et al., 2001, A&A, 365, 45
 Kessler M. F. et al., 1996, A&A, 315, L27
 Lehmann I. et al., 2000, A&A, 354, 35
 Lehmann I., Hasinger G., Murray S. S., Schmidt M., 2001, Proc. X-rays at Sharp Focus Chandra Science Symp.
 McHardy I. M. et al., 1998, MNRAS, 295, 641
 Morel T. et al., 2001, MNRAS, 327, 1187
 Moran E. C., Halpern J. P., Helfand D. J., 1996, ApJS, 106, 341
 Mushotzky R. F., Cowie L. L., Barger A. J., Arnaud K. A., 2000, Nat, 404, 459
 Nakanishi K., Akiyama M., Ohta K., Yamada T., 2000, ApJ, 534, 587
 Norman C. et al., 2001, ApJ, submitted (astro-ph/0103198)
 Oliver S. et al., 2000, MNRAS, 316, 749
 Read A. M., Ponman T. J., Strickland D. K., 1997, MNRAS, 286, 629
 Reeves J. N., Turner M. J. L., 2000, MNRAS, 316, 234
 Schartel N. et al., 1996, MNRAS, 283, 1015
 Schmidt M. et al., 1998, A&A, 329, 495
 Serjeant S. et al., 2000, MNRAS, 316, 768
 Smith D. A., Done C., 1996, MNRAS, 280, 355
 Stocke J. T., Case J., Donahue M., Shull J. M., Snow T. P., 1991, ApJ, 374, 72
 Turner T. J., George I. M., Nandra K., Mushotzky R. F., 1997, ApJ, 488, 164
 Veilleux S., Osterbrock D. E., 1987, ApSJ, 63, 295
 Veron P., Goncalves A. C., Veron-Cetty M.-P., 1997, A&A, 319, 52
 Voges W. et al., 1999, A&A, 349, 389
 Voges W. et al., 2000, IAU Circ. 7432, 3
 White N. E., Giommi P., Angelini L., 1994, A&AS, 185, 4111
 Willott C. J. et al., 2001, preprint (astro-ph/0105560)
 Xu C. et al., 1998, ApJ, 508, 576
 Zezas A. L., Georgantopoulos I., Ward M. J., 1998, MNRAS, 301, 915

This paper has been typeset from a $\text{\TeX}/\text{\LaTeX}$ file prepared by the author.