Extended search for supernovae-like neutrinos in NOvA coincident with LIGO/Virgo detections


This version is available from Sussex Research Online: http://sro.sussex.ac.uk/id/eprint/102529/

This document is made available in accordance with publisher policies and may differ from the published version or from the version of record. If you wish to cite this item you are advised to consult the publisher's version. Please see the URL above for details on accessing the published version.

Copyright and reuse:
Sussex Research Online is a digital repository of the research output of the University.

Copyright and all moral rights to the version of the paper presented here belong to the individual author(s) and/or other copyright owners. To the extent reasonable and practicable, the material made available in SRO has been checked for eligibility before being made available.

Copies of full text items generally can be reproduced, displayed or performed and given to third parties in any format or medium for personal research or study, educational, or not-for-profit purposes without prior permission or charge, provided that the authors, title and full bibliographic details are credited, a hyperlink and/or URL is given for the original metadata page and the content is not changed in any way.
Extended search for supernovallike neutrinos in NOvA coincident with LIGO/Virgo detections


(The NOvA Collaboration)
A search is performed for supernova-like neutrino interactions coincident with 76 gravitational wave events detected by the LIGO/Virgo Collaboration. For 40 of these events, full readout of the time around the gravitational wave is available from the NOvA Far Detector. For these events, we set limits on the fluence of the sum of all neutrino flavors of $F < 7(4) \times 10^{37} \text{cm}^{-2}$ at 90% C.L. assuming energy and time distributions corresponding to the Garching supernova models with masses $9.6(27) M_{\odot}$. Under the hypothesis that any given gravitational wave event was caused by a supernova, this corresponds to a distance of $r > 29(50) \text{kpc}$ at 90% C.L. Weaker limits are set for other gravitational wave events with partial Far Detector data and/or Near Detector data.

DOI: 10.1103/PhysRevD.104.063024

I. INTRODUCTION

Multimessenger astronomy is a rapidly expanding field, with exciting opportunities to simultaneously observe violent astrophysical events using gravitational waves (GWs), electromagnetic radiation, cosmic rays, and neutrinos. To date, a single gravitational wave event has been associated with electromagnetic activity [1–3], and none have been associated with the other channels. Not all gravitational waves and gravitational wave candidates to
date have been identified by the LIGO/Virgo Collaboration (LVC) with a particular production mechanism [4]. Although all clearly identified events are associated with compact object mergers, there remains the possibility that one or more were caused by a supernova, which are expected to produce gravitational waves, but with great uncertainty in predictions of the signal strength [5]. These potential supernovae may have evaded optical detection either because they were obscured by dust in the central Galaxy, or because they were “failed” supernovae in which the star collapsed, but did not explode [6].

In a previous paper [7] we described a broad search for signals, across the MeV to TeV range, associated with 26 gravitational wave events. We now focus on the possibility of detecting supernovalike neutrinos and present an improved search using the now-available larger catalog of gravitational wave events. The paper is organized as follows. In Sec. II, we introduce the NOvA detectors. Section III details the data set used in this analysis. Section IV explains how we simulate supernova neutrino interactions. Section V describes the improved selection of supernovalike neutrinos. Finally, Sec. VI gives the results.

II. DETECTORS

The NOvA experiment consists of two similar detectors, the Near Detector (ND) and the Far Detector (FD). The ND is located at the Fermi National Accelerator Laboratory (Fermilab), 100 m underground, while the FD is located near Ash River, Minnesota, on the surface with a modest overburden consisting of 1.25 m of concrete covered with 16 cm of barite gravel.

The NOvA detectors are segmented liquid scintillator tracking calorimeters. Alternating planes of cells are oriented horizontally and vertically, forming two views that can be used to reconstruct three-dimensional positions. The cells have a cross section of 4 cm by 6 cm and are 15.5 m (3.8 m) long in the FD (ND). The FD has 896 planes of cells and a total mass of 14 kt, whereas the ND has 214 planes and a total mass of 300 t. The last 20 planes at the north end of the ND are a muon catcher. They are interleaved with ten 10-cm-thick planes of steel for the purpose of measuring the energy of muons produced in beam interactions. The FD has no similar structure. The detectors are described in more detail elsewhere [8].

Light produced in the scintillator is collected by wavelength-shifting fibers and converted into electrical signals using avalanche photodiodes. These signals are continuously digitized at 2 MHz at the FD and 8 MHz at the ND. Samples rising above a threshold, called hits, are retained for further processing. Hits from all channels are collected into 50 μs blocks and can be saved for offline analysis if a software trigger requests them within about 20 minutes for the Far Detector and 30 minutes at the Near Detector. Triggers can either be based on the content of the data or on external signals. Two of the latter type of triggers are used in this analysis. First, when LVC publishes an observation of a gravitational wave candidate over the Gamma-ray Coordinates Network, we respond by reading out 45 s of continuous data from both the ND and FD, beginning 5.16 s prior to the gravitational wave time stamp. Second, we run a minimum bias pulser trigger on the FD which reads out 550 μs segments of data at a rate of 10 Hz. When only pulser data is available, we use a window of 1000 s centered on the gravitational wave time stamp to match the convention established by other neutrino observatories [9,10].

The NOvA detectors are exposed to Fermilab’s NuMI beam [11], a wideband neutrino beam with a peak at 2 GeV consisting mainly of either νμ or ¯νμ, depending on the operating mode. Typically, the beam is operated October through June with pulses of 10 μs separated by 1.3 s. For the purposes of the analysis reported here, the beam has no impact on the FD data since the number of beam neutrino interactions is negligible. However, it is a source of background at the ND; a procedure to remove beam backgrounds is detailed in Sec. V.

III. DATA SET

Tables I and II show a summary of NOvA data collected for each of the gravitational wave events and candidates (henceforth called “events”) announced by LVC to date in their two catalogs [4,12] and via the Gamma-ray Coordinates Network [13–42]. With the exception of four gravitational wave events, at least one of the NOvA detectors was operating and taking useful data for each event. LVC issued public triggers beginning with their “O3” run period in 2019; prior to that point, NOvA has only the FD pulser data. Thirteen events in O3 were only announced in the second LVC catalog and not via public trigger; we only have FD pulser data for these as well.

Of the remaining 52 GW events that did have public triggers, we recorded all or part of the desired 45 s of continuous data at the FD for 32 events, and at the ND for 40. In five cases, the ND recorded full readouts when the FD did not because it has a deeper data buffer. At each detector, data is read out approximately in time order; alerts that arrived when the data was near the end of the buffer resulted in partial readouts, as shown in the table. In the remaining three cases, the FD was down and the ND was up.

IV. SIMULATION

Supernova neutrino interactions are simulated for use in training the selector and for assessing signal significance. The simulation is based on the Garching 9.6 M⊙ and 27 M⊙ supernova flux models [43], with neutrino interactions produced with GENIE v3.0.6 [44], and the resulting particles tracked through the detector geometry using GEANT4 v10.4.2 [45]. The simulation only includes neutrinos above 10 MeV, with inverse beta decay on hydrogen (IBD)
and electron elastic scattering (ES) interactions included. Since NOvA is hydrocarbon-based, IBD strongly dominates over ES. IBD is the most important interaction for NOvA because it has a large cross section and produces a high-energy positron. The mean positron energy produced in the $^{9}\text{Be}$ simulation is $19.0(21.2)$ MeV.

In IBD interactions, both positrons and neutrons are simulated. Although NOvA is primarily sensitive to positrons and electrons, the 8 MeV of gammas from neutron capture on $^{35}\text{Cl}$ is also visible. The NOvA detectors are 16% chlorine by mass. After selection cuts, the FD has no significant sensitivity to electrons and positrons below 10 MeV; however, the ND is still marginally sensitive at this energy, so the simulation somewhat undercounts the neutrino interactions that would be selected in a real supernova.

### TABLE I. Summary of NOvA data taking during GW events [4,12–42] and 90% C.L. limits. The fluence limits on the two supernova models are in units of $10^{10}$ cm$^{-2}$. The distance limits are in kiloparsecs. When continuous data was read out in response to an LVC trigger, the number of seconds read is given for each detector. Otherwise ("untriggered"), pulser data is used in the case of the FD, and the ND is not used. In some cases one or both detectors were not running ("no data") and in two cases the FD was running, but not taking good data ("bad"). Events above the line have been considered by NOvA before; above and below the line events are arranged chronologically.

<table>
<thead>
<tr>
<th>Name</th>
<th>ND</th>
<th>FD</th>
<th>Fluence</th>
<th>Distance</th>
</tr>
</thead>
<tbody>
<tr>
<td></td>
<td></td>
<td></td>
<td>$\text{SN}_{27\odot}$</td>
<td>$\text{SN}_{9.6\odot}$</td>
</tr>
<tr>
<td>GW150914</td>
<td>Untriggered</td>
<td>Bad</td>
<td>...</td>
<td>...</td>
</tr>
<tr>
<td>GW151012</td>
<td>Untriggered</td>
<td>No data</td>
<td>...</td>
<td>...</td>
</tr>
<tr>
<td>GW151226</td>
<td>Untriggered</td>
<td>Untriggered</td>
<td>110</td>
<td>9</td>
</tr>
<tr>
<td>GW170104</td>
<td>Untriggered</td>
<td>Untriggered</td>
<td>300</td>
<td>500</td>
</tr>
<tr>
<td>GW170608</td>
<td>Untriggered</td>
<td>Untriggered</td>
<td>400</td>
<td>700</td>
</tr>
<tr>
<td>GW170729</td>
<td>Untriggered</td>
<td>Untriggered</td>
<td>240</td>
<td>400</td>
</tr>
<tr>
<td>GW170809</td>
<td>Untriggered</td>
<td>Untriggered</td>
<td>110</td>
<td>190</td>
</tr>
<tr>
<td>GW170814</td>
<td>Untriggered</td>
<td>Untriggered</td>
<td>120</td>
<td>200</td>
</tr>
<tr>
<td>GW170817</td>
<td>Untriggered</td>
<td>Untriggered</td>
<td>110</td>
<td>190</td>
</tr>
<tr>
<td>GW170818</td>
<td>Untriggered</td>
<td>Untriggered</td>
<td>180</td>
<td>330</td>
</tr>
<tr>
<td>GW170823</td>
<td>Untriggered</td>
<td>Untriggered</td>
<td>260</td>
<td>500</td>
</tr>
<tr>
<td>GW190408_181802</td>
<td>No data</td>
<td>No data</td>
<td>...</td>
<td>...</td>
</tr>
<tr>
<td>GW190412</td>
<td>Untriggered</td>
<td>Untriggered</td>
<td>170</td>
<td>280</td>
</tr>
<tr>
<td>GW190421_213856</td>
<td>Untriggered</td>
<td>Untriggered</td>
<td>210</td>
<td>400</td>
</tr>
<tr>
<td>GW190425</td>
<td>Untriggered</td>
<td>Untriggered</td>
<td>120</td>
<td>190</td>
</tr>
<tr>
<td>GW190426_152155</td>
<td>44.7 s</td>
<td>Untriggered</td>
<td>13</td>
<td>19</td>
</tr>
<tr>
<td>GW190503_185404</td>
<td>Untriggered</td>
<td>Untriggered</td>
<td>150</td>
<td>270</td>
</tr>
<tr>
<td>S190510g</td>
<td>Untriggered</td>
<td>Untriggered</td>
<td>170</td>
<td>280</td>
</tr>
<tr>
<td>GW190512_180714</td>
<td>Untriggered</td>
<td>Untriggered</td>
<td>190</td>
<td>330</td>
</tr>
<tr>
<td>GW190513_205428</td>
<td>24.7 s</td>
<td>Untriggered</td>
<td>14</td>
<td>20</td>
</tr>
<tr>
<td>GW190517_055101</td>
<td>Untriggered</td>
<td>Untriggered</td>
<td>120</td>
<td>200</td>
</tr>
<tr>
<td>GW190519_153544</td>
<td>Untriggered</td>
<td>Untriggered</td>
<td>140</td>
<td>250</td>
</tr>
<tr>
<td>GW190521</td>
<td>45.0 s</td>
<td>45.0 s</td>
<td>6</td>
<td>10</td>
</tr>
<tr>
<td>GW190521_074359</td>
<td>Untriggered</td>
<td>Untriggered</td>
<td>170</td>
<td>280</td>
</tr>
<tr>
<td>GW190602_175927</td>
<td>45.0 s</td>
<td>45.0 s</td>
<td>6</td>
<td>12</td>
</tr>
<tr>
<td>GW190630_185205</td>
<td>45.0 s</td>
<td>45.0 s</td>
<td>5</td>
<td>9</td>
</tr>
<tr>
<td>GW190701_203306</td>
<td>45.0 s</td>
<td>45.0 s</td>
<td>6</td>
<td>11</td>
</tr>
<tr>
<td>GW190706_222641</td>
<td>45.0 s</td>
<td>17.5 s</td>
<td>2.5</td>
<td>5</td>
</tr>
<tr>
<td>GW190707_093326</td>
<td>Untriggered</td>
<td>Untriggered</td>
<td>220</td>
<td>400</td>
</tr>
<tr>
<td>GW190913_052954</td>
<td>Untriggered</td>
<td>Untriggered</td>
<td>170</td>
<td>280</td>
</tr>
<tr>
<td>GW190913_134308</td>
<td>Untriggered</td>
<td>Untriggered</td>
<td>160</td>
<td>270</td>
</tr>
<tr>
<td>GW190924_180646</td>
<td>Untriggered</td>
<td>Untriggered</td>
<td>140</td>
<td>240</td>
</tr>
<tr>
<td>GW190954_065416</td>
<td>Untriggered</td>
<td>Untriggered</td>
<td>280</td>
<td>500</td>
</tr>
<tr>
<td>GW190957_092055</td>
<td>Untriggered</td>
<td>Untriggered</td>
<td>140</td>
<td>240</td>
</tr>
<tr>
<td>GW1909620_030421</td>
<td>Untriggered</td>
<td>Untriggered</td>
<td>270</td>
<td>400</td>
</tr>
<tr>
<td>GW190708_232457</td>
<td>Untriggered</td>
<td>Untriggered</td>
<td>150</td>
<td>270</td>
</tr>
<tr>
<td>S190718g</td>
<td>18.3 s</td>
<td>Untriggered</td>
<td>17</td>
<td>23</td>
</tr>
<tr>
<td>GW190719_215514</td>
<td>Untriggered</td>
<td>Bad</td>
<td>...</td>
<td>...</td>
</tr>
<tr>
<td>GW190720_000836</td>
<td>45.0 s</td>
<td>45.0 s</td>
<td>4</td>
<td>6</td>
</tr>
<tr>
<td>GW190727_060333</td>
<td>45.0 s</td>
<td>45.0 s</td>
<td>5</td>
<td>9</td>
</tr>
</tbody>
</table>
Besides undercounting low-energy neutrino interactions, the simulation also does not include various interaction channels on carbon such as $\nu_e + ^{12}C \rightarrow e^- + ^{12}N$, nor similar channels involving other isotopes in the NOvA materials, although in many cases, these interactions would be easily visible. The limits set below are therefore conservative, although IBD would dominate over these other channels even if they were included. We use a model without neutrino oscillations or other flavor-changing effects because there is not enough information available to know whether these effects would increase or decrease the number of neutrinos observed by NOvA [43].

V. ANALYSIS

Relative to our previous report [7], the clustering algorithm for grouping hits into supernova neutrino event candidates has been greatly improved. Previously, a cluster was defined as a pair of hits with one hit in each view. Now a cluster may have two to seven hits associated in time and space. Clusters of greater than seven hits are rejected as being too large to have been produced by a supernova neutrino interaction. In the ND, we allow clusters with all hits in a single view. However, at the FD, three-dimensional position information is essential for reducing background, so clusters must include hits in both views. Similarly, ND
clusters may be noncontiguous, with gaps either between hits within a detector plane or between detector planes, but FD clusters must be contiguous to reduce background. Previously, we excluded the muon catcher region of the ND; clusters in this region are now accepted.

Critical to the reduction of background, particularly at the FD, is the inclusion of several new variables in the classifier that relate the distance in time and space between candidate hit clusters and recent cosmic rays. Michel electrons from stopping muons are a common background in the FD, occurring at a rate of 40 kHz. Most Michel electrons are identified by close association with track ends, but a small fraction of apparent Michel electrons appear far from the track end, either because of reconstruction failures, inefficiencies in producing hits, complex particle interactions, or some combination of these. Candidate clusters are judged based on their proximity to the track end, to any point along the track, as well as to any hit in a large cluster of activity with no reconstructed tracks.

Supernova neutrino-like hit clusters are separated into signal and background samples using the scikit-learn\cite{scikit-learn} package’s RandomForestClassifier class. The classifier is trained with simulated $9.6 \ M_\odot$ supernova interactions and real minimum-bias data from the NOvA detectors. The classifier was optimized separately for the ND and the FD. Further, it is optimized separately for the two cases of FD data: continuous readout and pulser. The pulser data must be treated differently because the lookback time for cosmic rays that may have produced a background cluster is reduced. Additionally, since the live time is smaller, efficiency is prioritized over background reduction. In all three cases, the figure of merit\cite{figure_ofMerit}

$$a = \frac{\text{signal}}{2 + \sqrt{\text{background}}}$$

is optimized, with $a = 1.292$ to optimize 90% C.L. limits. The resulting efficiencies for IBD positrons are shown in Fig. 1. Efficiencies for ES electrons, as a function of electron energy, are very similar. At the ND, neutron captures from IBD are selected with 2% efficiency, while at the FD the neutron capture efficiency is negligible for purposes of the signal; only 0.02%. No attempt is made in the analysis to associate positron and neutron delayed coincidences in either detector.

Compared with our previous analysis method, the rate of selected background candidates in the FD, for continuous readout, is reduced by a factor of 80, from 460 to 6 Hz, while the signal efficiency for IBD positrons is reduced from 7.8% to 4.3%. The reduced signal efficiency is a consequence of the optimization described above. In the previous analysis, the same selection was used for FD continuous-readout and FD pulser data. In this analysis, the pulser background rate is reduced to 55 Hz, while the signal efficiency is increased to 9.0%, or 0.3 Hz and 0.05% taking into account the 0.55% live time. In the ND, the rate of selected background candidates has been slightly reduced from 0.5 to 0.4 Hz while the signal efficiency has been increased from 12% to 44%.

Since the neutrino event classifier is trained on real detector data, no explicit identification of the background components is made. The FD background likely contains significant components from cosmogenic thermal neutron

![Efficiency plot](image-url)
captures, cosmogenic $^{12}$B and $^{13}$N beta decays, and single-hit uranium/thorium-chain radioactivity paired with unrelated single-hit electronics noise. The latter is possibly a significant component of the ND background as well, but cosmogenic activity is strongly suppressed compared to the FD.

For 16 of the 40 GW events with ND data, the NuMI beam was in operation. Data at the ND is taken in 5 ms segments. Any 5 ms data segment is rejected if it overlaps with a beam pulse or the time up to 3 ms following a beam pulse. This conservative cut removes all prompt beam activity, muon decays, and neutron captures from thermal neutrons that were produced in the detector and remained in the detector until captured. Some neutron captures can be delayed up to several milliseconds if thermal neutrons spend time in the air surrounding the detector; the 3 ms cut rejects a large majority of these neutrons.

For each gravitational wave event, we first examine the selected clusters in 1-second bins searching for any significant excess over background, where the background level is determined in situ from the 45 s readout (or 1000 s window in the case of FD pulser data). Second, we assume that a supernova burst begins at the gravitational wave time stamp and set limits on its strength for the case of the Garching 9.6 $M_\odot$ and 27 $M_\odot$ models. Because NOvA’s efficiency rises rapidly with neutrino energy between 10 and 30 MeV, the higher neutrino energies in the 27 $M_\odot$ model result in stronger fluence limits.

Depending on the state of the two NOvA detectors and whether a trigger was received from LVC, several different types of data sets can be available. The best case is when a timely trigger was received and we read 45 s of continuous data from the ND and FD. In this case, a joint analysis is done using the data from the two detectors. The FD provides more statistical power, but the ND still makes a significant contribution. In some cases, continuous data is available from the ND, but only pulser data from the FD. Again, a joint analysis is performed, but in this case, the ND provides nearly all the statistical power. In some cases, the continuous data from ND or FD is not a complete 45 s, but in all those cases enough was read out to establish the background level and allow the analysis to be run without modification. The background level is not determined with as much precision in these cases, leading to a slight weakening of limits. Finally, in some cases, data from only one detector is available. The status for all GW events is shown in Table I.

VI. RESULTS

No excess over background is observed for any gravitational wave event at any time within the analyzed window. Background rates were stable at both detectors, being around 5 Hz at the FD for the continuous-readout selection, 0.3 Hz for the FD pulser selection and 0.4 Hz at the ND. Assuming all selected clusters are background, the limits depend on statistical fluctuations in the background

in the first few seconds after the gravitational wave time stamp. A typical event is shown in Fig. 2.

For each GW event, 90% C.L. limits are set on the fluence of the sum of all neutrino flavors, $F$, under the assumption of the two Garching supernova models discussed above, without flavor-changing effects. The limits are set via a fit to the time series of neutrino candidates with two parameters: the background rate and the signal strength, with the signal templates as shown in Fig. 2. A Bayesian approach is used with flat priors in each parameter.

A posterior probability density function, profiled over the background level, is constructed by scanning over the background-only fit, normalized to 10 kpc. The number of neutrino candidates per second is corrected for live time, which is slightly under 100% in the ND because of beam removal, and in the final bins because readout ends at 39.86 s. The limits set are weaker than the median case because of a slight excess in the 0–5 s bins in the statistically dominant FD.

$$- \log L = \sum_i \left( m_i - d_i + d_i \log \frac{d_i}{m_i} \right)$$

is computed for the background normalization that minimizes $- \log L$, where $m_i$ is the number of events predicted

FIG. 2. A typical GW event with both FD (top) and ND (bottom) continuous readout, S200213t. The two supernova models are shown, normalized to 10 kpc. The number of neutrino candidates per second is corrected for live time, which is slightly under 100% in the ND because of beam removal, and in the final bins because readout ends at 39.86 s. The limits set are weaker than the median case because of a slight excess in the 0–5 s bins in the statistically dominant FD.
by the model in bin \(i\) and \(d_i\) is the number of events observed in bin \(i\). The probability density is proportional to \(L\). The resulting curve is integrated numerically up to 90% of the total, and this sets the 90% upper limit on signal strength, \(s_{90}\). Because the signal would decrease as \(1/r^2\), where \(r\) is the distance to the hypothetical supernova, the 90% lower limit on distance is \(d_{90} = 10\, kpc/\sqrt{s_{90}}\). Given the number of neutrinos predicted by each model, \(N = 6.8(11) \times 10^{27}\) for 9.6(27) \(M_\odot\), fluence limits, \(F_{90}\), are related to the distance limits via

\[
F_{90} = \frac{N}{4\pi d_{90}^2}.
\]

No systematic effects are explicitly included in the procedure, but as detailed in Sec. IV, our estimate of the rate of detectable neutrino interactions is conservative; we believe this conservatism is sufficient to cover any systematic effects in signal efficiency. All limits are shown in Table I and a discussion of notable features thereof follows.

For gravitational wave events in which we read out continuous FD data in response to an LVC trigger (the best case), fluence limits range between \(F < 4 \times 10^{10} \text{ cm}^{-2}\) and \(F < 15 \times 10^{10} \text{ cm}^{-2}\), assuming the 9.6 \(M_\odot\) Garching model. In this model, 22% of the neutrinos are \(\bar{\nu}_e\), to which NOvA is primarily sensitive. The median limit is \(7 \times 10^{10} \text{ cm}^{-2}\). Similarly, for the 27 \(M_\odot\) model, in which 23% of the flux is \(\bar{\nu}_e\), we set limits ranging from \(F < 2.1 \times 10^{10} \text{ cm}^{-2}\) to \(F < 9 \times 10^{10} \text{ cm}^{-2}\) with a median of \(4 \times 10^{10} \text{ cm}^{-2}\). If interpreted as limits on the distance to a hypothetical supernova, we exclude a 9.6 \(M_\odot\) supernova in the median case, at 90% C.L., closer than 29 kpc. For the event with the strongest exclusion, S200115j, we exclude a 9.6 \(M_\odot\) supernova closer than 40 kpc. For the 27 \(M_\odot\) model, we exclude a supernova, in the median case, closer than 50 kpc, and for S200115j, 70 kpc.

In the next best case, we have continuous ND data, but either have no FD data or only pulser data from the FD data. In the latter case, the limit is strongly dominated by the ND data. Fluence limits for the 9.6 \(M_\odot\) model range from \(F < 19 \times 10^{10} \text{ cm}^{-2}\) to \(F < 29 \times 10^{10} \text{ cm}^{-2}\), and for the 27 \(M_\odot\) from \(F < 13 \times 10^{10} \text{ cm}^{-2}\) to \(F < 22 \times 10^{10} \text{ cm}^{-2}\). Because of the ND’s lower background, the efficiency for selecting lower-energy neutrinos is higher than the FD. The flux model therefore has less effect on fluence limits dominated by ND data. The median distance limit for a 9.6(27) \(M_\odot\) supernova is 16(24) kpc.

Finally, when using only FD pulser data, fluence limits range from \(F < 190 \times 10^{10} \text{ cm}^{-2}\) to \(F < 700 \times 10^{10} \text{ cm}^{-2}\) for the 9.6 \(M_\odot\) model and from \(F < 110 \times 10^{10} \text{ cm}^{-2}\) to \(F < 400 \times 10^{10} \text{ cm}^{-2}\) for the 27 \(M_\odot\) model. Even with only FD pulser data, some exclusion of supernovae in or behind the Galactic core (at \(\sim\)8 kpc), whose optical signal may have been obscured, is possible, with distance limits ranging from 2.9–6 kpc for the 9.6 \(M_\odot\) case and 5–9 kpc for the 27 \(M_\odot\) case.

The 26 GW events analyzed in our previous report are reanalyzed using the improved analysis. The limits quoted for the seven previously analyzed events that include FD and/or ND continuous readout are now stronger, in the median case, by a factor of 3, and in no case is the result we now give weaker than our previously published result. However, for events with only FD pulser data, the new analysis techniques only yield a 40% improvement in fluence limits. There are four GW events that, in the new analysis, have a weaker limit for at least one of the two supernova models: GW170608, GW170729, GW170823 and GW190521_074359. This is an expected consequence of using an analysis that is almost entirely different than our previous analysis, such that there is little correlation between the hits selected previously and now.

**VII. CONCLUSIONS**

We have searched for supernova-like neutrinos coincident with 76 gravitational wave events reported by LVC. No excess consistent with such neutrinos was found. Assuming a burst of supernova-like neutrinos beginning at LVC’s reconstructed gravitational wave time, we set limits on the fluence of supernova-like neutrinos under two supernova models. In the 32 cases with full FD data, these limits are sufficient to largely exclude the possibility that any of the gravitational waves originated from a stellar core collapse in our Galaxy. This includes the “failed supernovae” in which there is no explosion and/or scenarios that lead to early black hole formation, since similar neutrino luminosities are expected in any of these cases [48,49]. Our search complements those performed by other neutrino observatories [10,50–56]. The NOvA detectors will continue to operate for several years, including during the upcoming O4 run of LVC.

**ACKNOWLEDGMENTS**

This document was prepared by the NOvA Collaboration using the resources of the Fermi National Accelerator Laboratory (Fermilab), a U.S. Department of Energy, Office of Science, HEP User Facility. Fermilab is managed by Fermi Research Alliance, LLC (FRA), acting under Contract No. DE-AC02-07CH11359. This work was supported by the U.S. Department of Energy; the U.S. National Science Foundation; the Department of Science and Technology, India; the European Research Council; the MSMT CR, GA UK, Czech Republic; the RAS, RFBR, the Ministry of Education and Science of the Russian Federation (RMES), RSF, and BASIS Foundation, Russia; CNPq and FAPEG, Brazil; STFC, UKRI, and the Royal Society, United Kingdom; and the State and University of Minnesota. We are grateful for the contributions of the staffs of the University of Minnesota at the Ash River Laboratory and of Fermilab.


[34] LIGO Scientific Collaboration and Virgo Collaboration, S200129m, GCN 26926 (2020).


[54] K. Abe et al. (Super-Kamiokande Collaboration), Search for neutrinos in coincidence with gravitational wave events from the LIGO-Virgo O3a observing run with the Super-Kamiokande detector, arXiv:2104.09196.
